

Star Formation in Cooling Winds of Massive Star Clusters

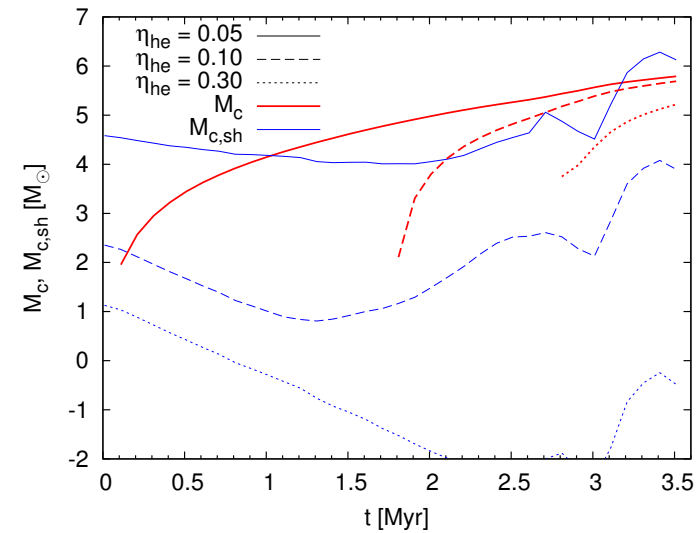
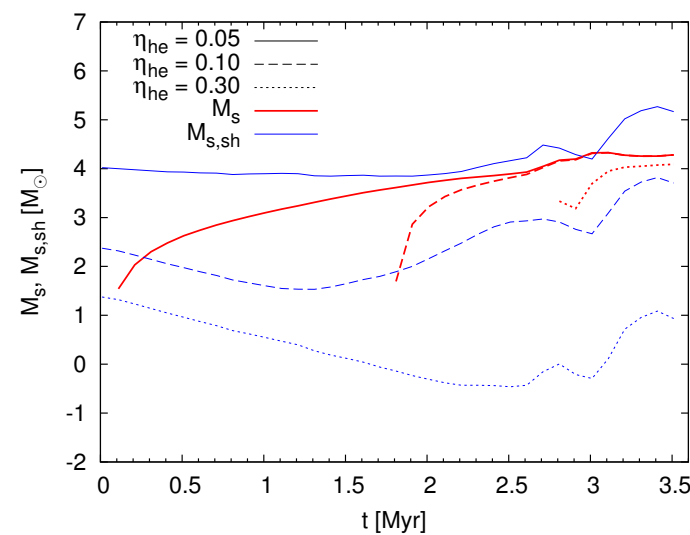
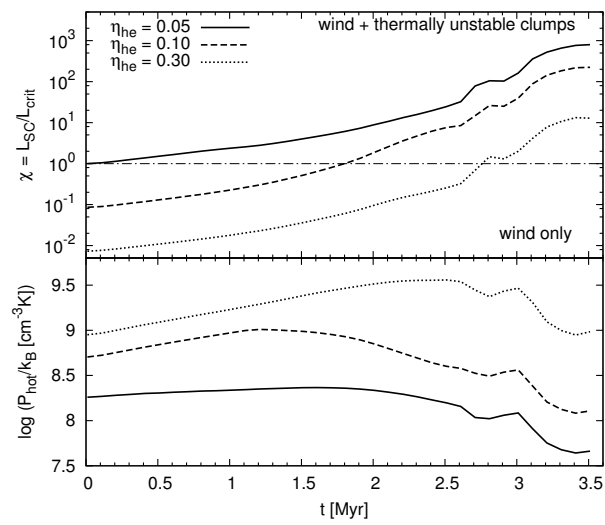
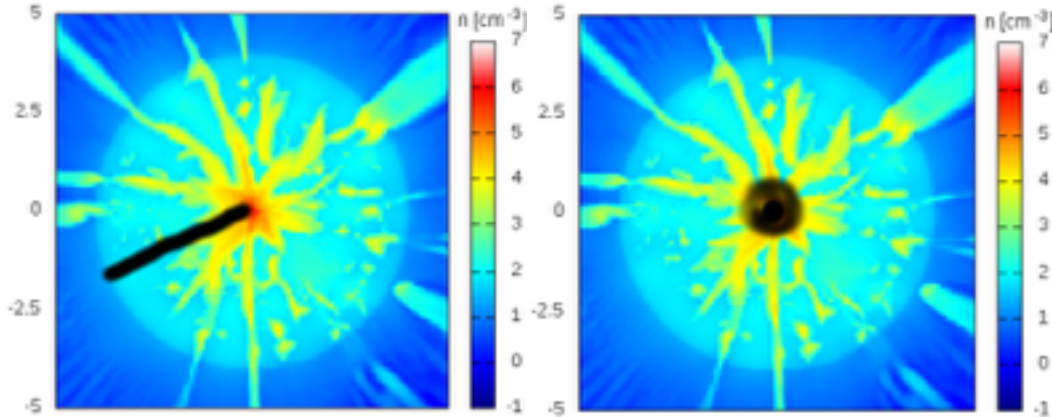
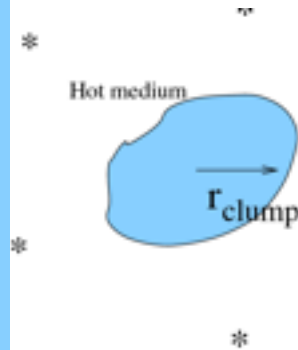
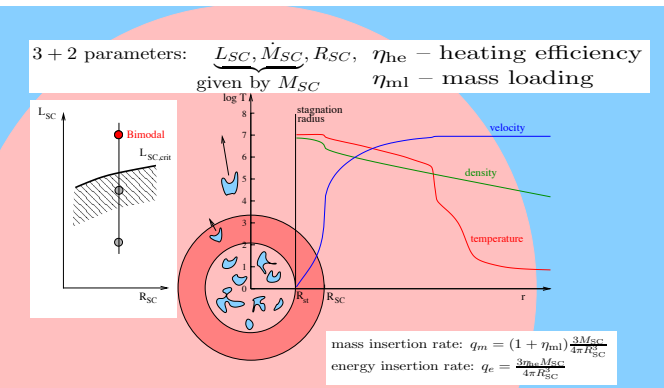
Jan Palouš

Soňa Ehlerová, Richard Wünsch

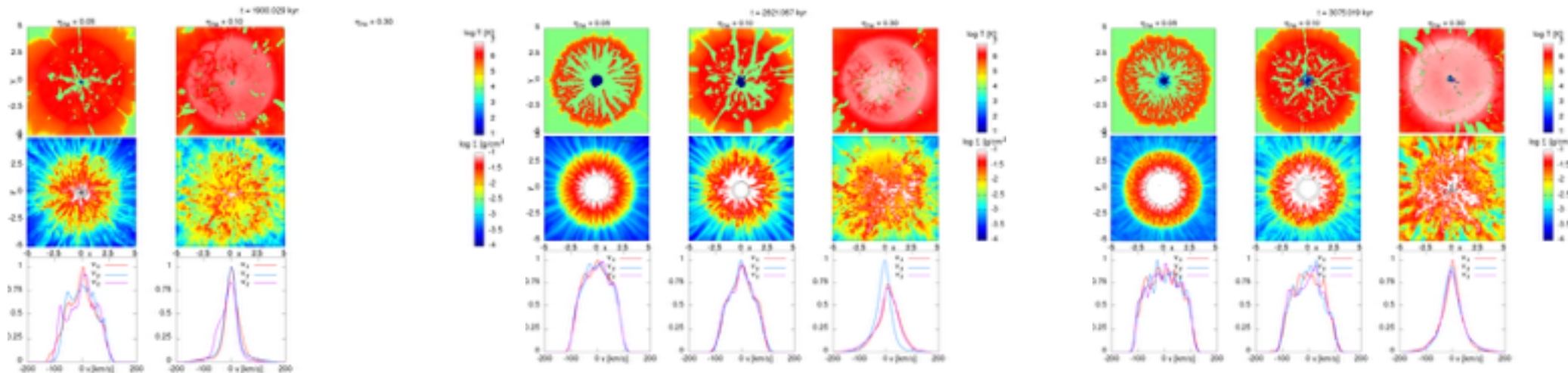
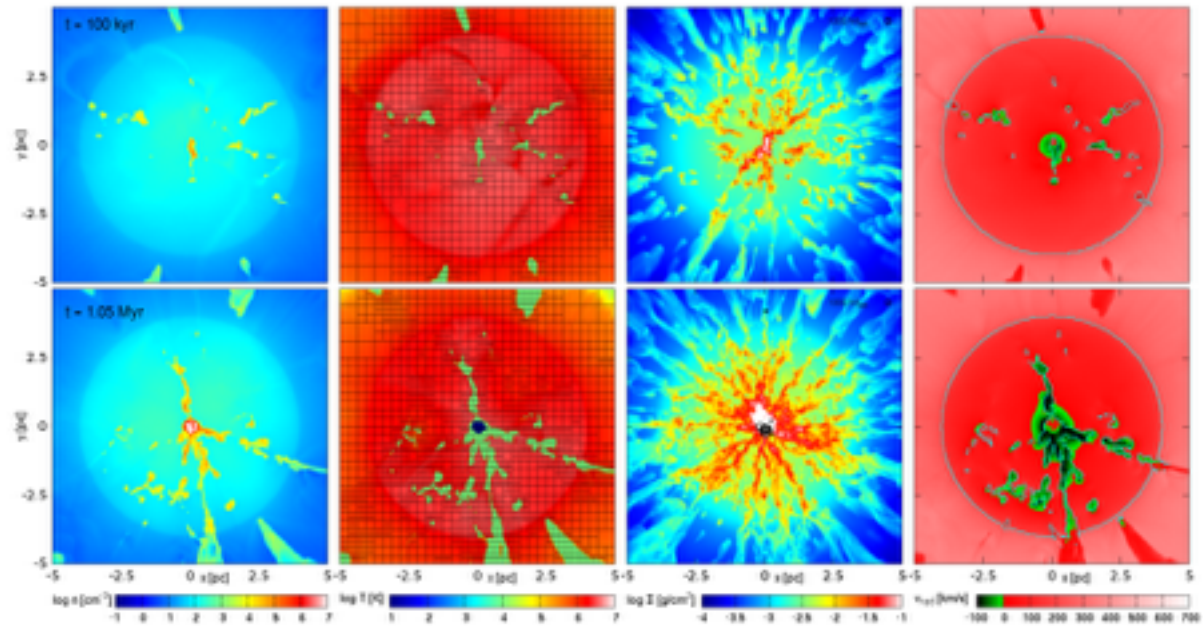
Guillermo Tenorio-Tagle

Astronomical Institute, CAS ,Czech Republic
Instituto Nacional de Astrofísica Optica y
Electronica, Puebla, Mexico

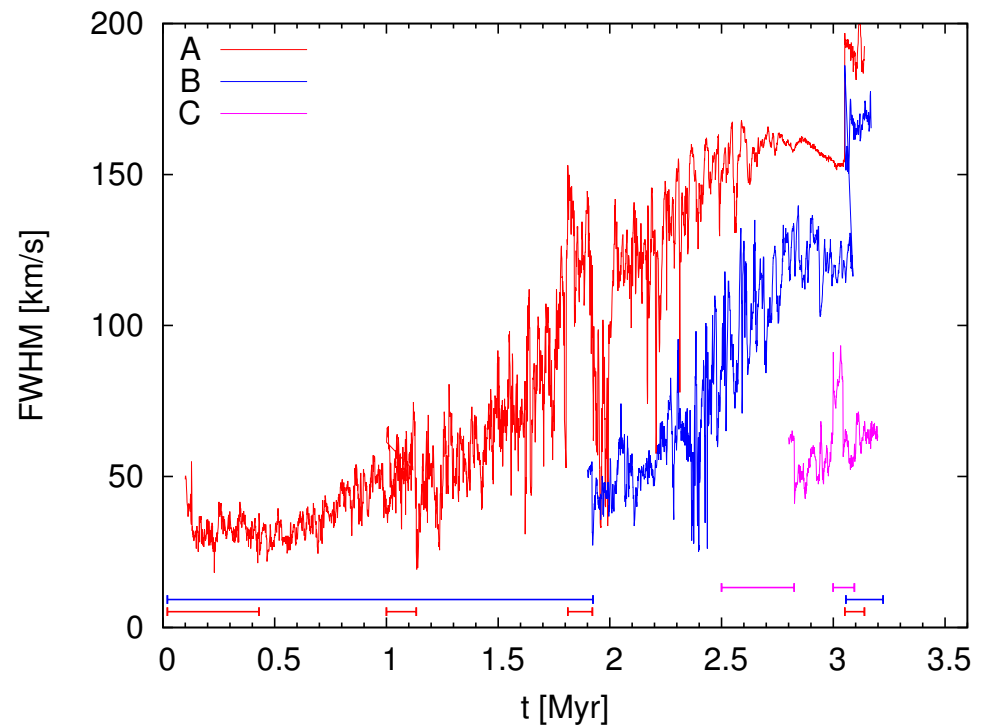
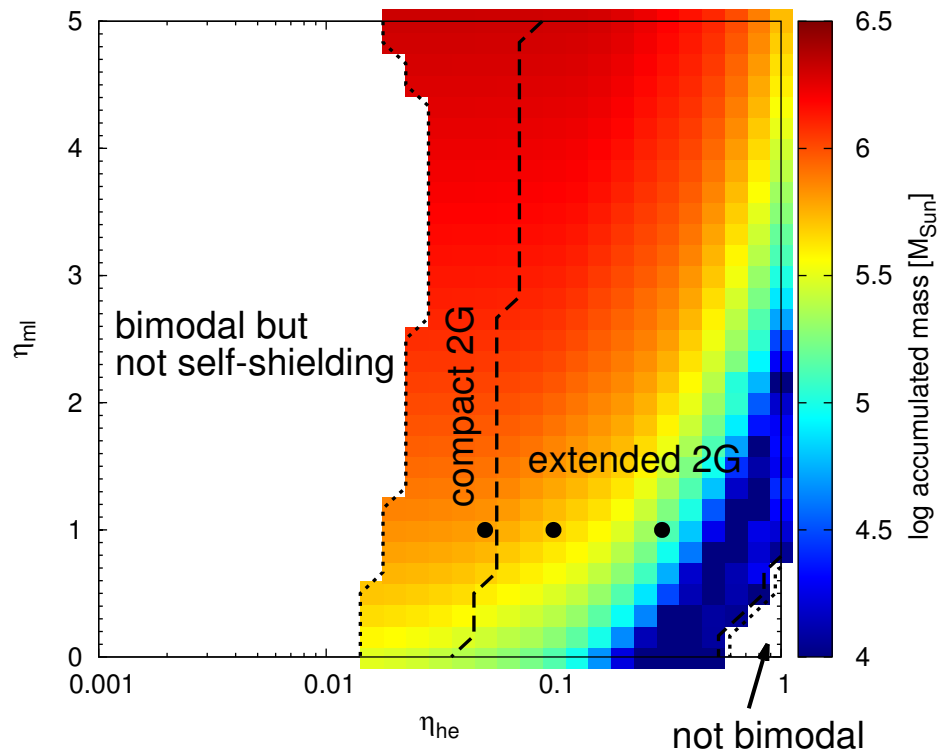
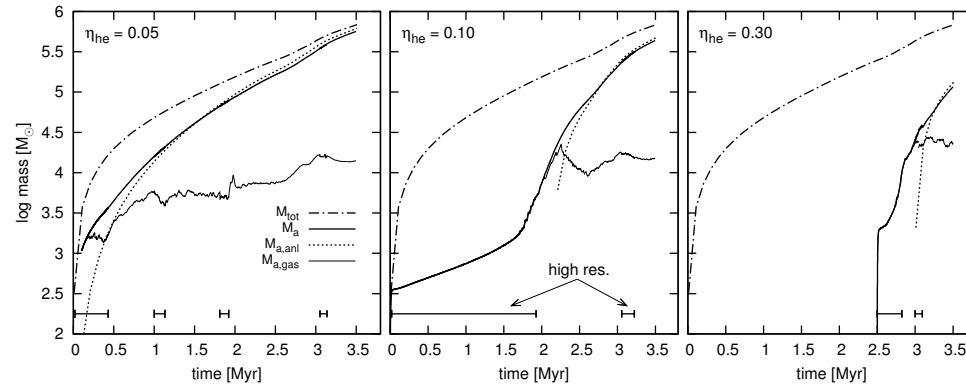
Cooling Winds Streams or Central Clump



Cooling Winds of Massive Star Clusters



Cooling Winds of Massive Star Clusters



Conclusions

- Secondary star formation within first 3.5 Myr;
- $MG1 = 10^7 \text{ Msun}$
 $0.01 < \text{heating efficiency} < 0.3$,
 $0.5 < \text{mass loading} < 2$,
 $2 \cdot 10^5 < MG2 < 10^6 \text{ Msun}$
 $1.5 < t_{\text{self-shield}} < 2.5 \text{ Myr}$;
- Compact cluster of the second generation;
- Removal of the large part of the first generation by tides;
- Line profiles: predicted versus observed